

Traversable wormholes via a double trace deformation

Ping Gao,^a Daniel Louis Jafferis^a and Aron C. Wall^b

^a*Center for the Fundamental Laws of Nature, Harvard University,
Cambridge, MA, U.S.A.*

^b*School of Natural Sciences, Institute for Advanced Study,
Princeton, NJ, U.S.A.*

E-mail: pgao@g.harvard.edu, jafferis@physics.harvard.edu,
aroncwall@gmail.com

ABSTRACT: After turning on an interaction that couples the two boundaries of an eternal BTZ black hole, we find a quantum matter stress tensor with negative average null energy, whose gravitational backreaction renders the Einstein-Rosen bridge traversable. Such a traversable wormhole has an interesting interpretation in the context of ER=EPR, which we suggest might be related to quantum teleportation. However, it cannot be used to violate causality. We also discuss the implications for the energy and holographic entropy in the dual CFT description.

KEYWORDS: Black Holes, Gauge-gravity correspondence

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1 Introduction

Traversable wormholes have long been a source of fascination as a method of long distance transportation [38]. However, such configurations require matter that violates the null energy condition, which is believed to apply in physically reasonable classical theories. In quantum field theory, the null energy condition is false, but in many situations there are other no-go theorems that rule out traversable wormholes.

In this work we find that adding certain interactions that couple the two boundaries of eternal AdS-Schwarzschild results in a quantum matter stress tensor with negative average null energy, rendering the wormhole traversable after gravitational backreaction. The coupling we turn on has the effect of modifying the boundary conditions of a scalar field in the bulk, which changes the metric at 1-loop order.

Violation of the averaged null energy condition (ANEC) is a prerequisite for all traversable wormholes [25, 39, 52, 53]. It states that there must be infinite null geodesics passing through the wormhole, with tangent vector k^μ and affine parameter λ , along which

$$\int_{-\infty}^{+\infty} T_{\mu\nu} k^\mu k^\nu d\lambda < 0. \quad (1.1)$$

The physical picture is that by Raychaudhuri's equation for null geodesic congruence, light rays will defocus only when ANEC is violated. In that case, the light rays that focus in one end of the wormhole can defocus when going out the other end.

There are reasonable arguments that the ANEC is always obeyed along infinite achronal geodesics [22, 31, 33, 34, 55].¹ This is sufficient to rule out traversable wormholes joining two otherwise disconnected regions of spacetime [22]. Furthermore, the generalized second law

¹A set of points is achronal if no two of the points can be connected by a timelike curve; otherwise it is *chronal*.

(GSL) of causal horizons also rules out traversable wormholes connecting two disconnected (asymptotically flat or AdS) regions, due to the fact that the future horizon of a lightray crossing through the wormhole has divergent area at very early times, which contradicts the increase of generalized entropy along the future horizon [57].

For small semiclassical perturbations to a stationary causal horizon, both the GSL and the ANEC follow from lightfront quantization methods that are valid for free or super-renormalizable field theories [56]. (There is also evidence that these results extend to more general field theories [13, 19, 26, 27, 32]).

In our configuration, signals from early times on the horizon can intersect it again at late times, by passing through the directly coupled boundaries. The causal structure of the manifold is modified as a result, changing the commutation relations along null geodesics through the wormhole and making them no longer achronal. For the same reason, a causal horizon extending through the wormhole intersects itself, removing the piece with divergent area. Hence the above impossibility results do not apply. The negative energy matter in our configuration is similar to the Casimir effect, since the interaction between the boundaries implies that the radial direction is effectively a compact circle.

Another problematic aspect of traversable wormholes is that they have the potential to lead to causal inconsistencies. For example, by applying a boost to one end of a wormhole one could attempt to create a configuration with closed time-like curves [39]. The direct interaction of the boundaries that we require implies that no such paradoxes may arise (for a more detailed discussion, see section 4).

The traversable wormhole we find is the first such solution that has been shown to be embeddable in a UV complete theory of gravity. A phenomenological model of a static BTZ wormhole that becomes traversable as a result of nonperturbative effects in a $1/c$ expansion was proposed in [47] (c being the central charge), however it was not shown that the metric obeys any field equations. A traversable wormhole solution of five dimensional Einstein-Gauss-Bonnet gravity was found in [3, 9, 50], however that low energy effective theory appears to lack a UV completion [14]. Another example was found [6] in a theory with a conformally coupled scalar, in a regime in which the effective Newton's constant becomes negative. This suggests that this solution also cannot arise in a UV complete model. The important fact that the boundary CFT dual of a traversable wormhole must involve interactions between the two CFTs was noted in [3, 47].

The eternal black hole with two asymptotically AdS regions is the simplest setting to investigate these questions [36]. We will deform the system by turning on a relevant double trace deformation [1]

$$\delta S = \int dt d^{d-1}x h(t, x) \mathcal{O}_R(t, x) \mathcal{O}_L(-t, x), \tag{1.2}$$

where \mathcal{O} is a scalar operator of dimension less than $d/2$, dual to a scalar field φ . This connects the boundaries with the same time orientation, since the t coordinate runs in opposite directions in two wedges (see figure 1a). The small deformation $h(t, x)$ has support only after some turn-on time t_0 . By the AdS/CFT correspondence, we can be certain that this relevant deformation corresponds to a consistent configuration in quantum gravity.

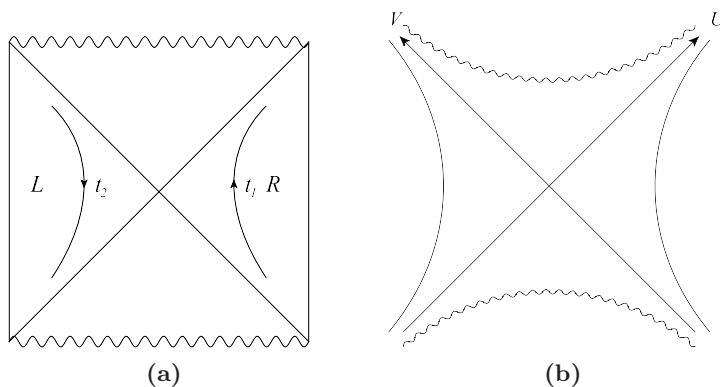


Figure 1. (a) is the Penrose diagram and (b) shows the Kruskal coordinates of the eternal black hole.

The eternal black hole has a Killing symmetry which is time-like outside the horizon. Null rays along the horizon $V = 0$ pass through the bifurcation surface of the Killing vector, and asymptote to $t \rightarrow -\infty$ on the left boundary and $t \rightarrow +\infty$ on the right boundary (see figure 1b). Denote the affine parameter along this ray as U . In the linearized analysis around this solution, the throat will become marginally traversable if $\int dU T_{UU} < 0$, where the integral is over the whole U coordinate.

It is instructive to see explicitly in this case that a small spherically symmetric perturbation of the stress tensor $T_{\mu\nu} \sim O(\epsilon)$ results in a traversable wormhole exactly when the ANEC is violated, by solving the linearized Einstein equation for $h_{\mu\nu} = \delta g_{\mu\nu} \sim O(\epsilon)$. Using Kruskal coordinates for the background metric, we find that at $V = 0$,

$$\frac{(d-2)}{4} [((d-3)r_h^{-2} + (d-1)\ell^{-2})(h_{UU} + \partial_U(Uh_{UU})) - 2r_h^{-2}\partial_U^2 h_{\phi\phi}] = 8\pi G_N T_{UU} \quad (1.3)$$

where ϕ is the azimuthal angle, r_h is the horizon radius of the black hole and the cosmological constant is $\Lambda = -\frac{(d-2)(d-1)}{2l^2} < 0$.

Since the deformation of the Hamiltonian is small, after the scrambling time, the fields ought to approach a stationary state with respect to an asymptotic Killing symmetry $U\partial U$. Hence T_{UU} must decay faster than U^{-2} , as does each term in l.h.s. of (1.3) after imposing a suitable gauge at past and future infinity. Therefore, if we integrate (1.3) over U the total derivative terms drop out and we obtain

$$8\pi G_N \int dU T_{UU} = \frac{(d-2)}{4} ((d-3)r_h^{-2} + (d-1)\ell^{-2}) \int dU h_{UU} \quad (1.4)$$

Linearized diffeomorphisms around the stationary black hole background act on $h_{\mu\nu}$, but when the AdS asymptotic conditions are imposed the quantity $\int_{-\infty}^{+\infty} dU h_{UU}$ is gauge invariant. Note that the null ray originating on the past horizon is given in coordinates by

$$V(U) = -(2g_{UV}(0))^{-1} \int_{-\infty}^U dU h_{UU} \quad (1.5)$$

after including the perturbation to linear order, where $g_{UV}(0) < 0$ is the UV component of the original metric on the $V = 0$ slice. If the ANEC is violated, $V(+\infty) < 0$, and a light ray from left boundary will hit the right boundary after finite time.

Note that if there existed any state in which the wormhole was traversable in the system defined by the decoupled Hamiltonian, $H_L + H_R$, then it would contradict the AdS/CFT duality. This is because in the decoupled system, no operator on the left can influence the right, which implies that no signal can be transmitted between the boundaries through the bulk.

At the linearized level, if one modifies the state as $|\text{tfd}\rangle \rightarrow e^{i\epsilon A}|\text{tfd}\rangle$ for small ϵ , the average null energy becomes $\langle \int dUT_{UU} \rangle = i\epsilon \langle [\int dUT_{UU}, A] \rangle$. If this were non-vanishing for any operator A , then by adjusting the sign of ϵ , the throat could be made traversable. It is easy to check that the expectation value of this commutator indeed vanishes.

In fact, $|\text{tfd}\rangle$ is invariant under $H_R - H_L$, which corresponds to the bulk Killing symmetry $i\partial_t$ (note the directions are opposite in left and right wedges). On the horizon $V = 0$, one can show $\partial_t = U\partial_U$ in Kruskal coordinates, which is just a dilation of the U direction. Note that under the $U \rightarrow \lambda U$ scaling, $T_{UU} \rightarrow \lambda^{-2}T_{UU}$ and $dU \rightarrow \lambda dU$, which implies $[H_R - H_L, \int dUT_{UU}] = -i \int dUT_{UU}$. Therefore

$$\begin{aligned} (H_R - H_L) \int dUT_{UU}|\text{tfd}\rangle &= \left[H_R - H_L, \int dUT_{UU} \right]|\text{tfd}\rangle + \int dUT_{UU}(H_R - H_L)|\text{tfd}\rangle \\ &= -i \int dUT_{UU}|\text{tfd}\rangle. \end{aligned} \tag{1.6}$$

This implies that $\int dUT_{UU}|\text{tfd}\rangle$ is either an eigenvector of $H_R - H_L$ with eigenvalue $-i$, or identically zero. Since $H_R - H_L$ is a Hermitian operator, whose eigenvalues must be real, it follows that $\int dUT_{UU}|\text{tfd}\rangle = 0$. In other words, T_{UU} in the modified state along $U > 0$ will exactly cancel that along $U < 0$. Beyond the linearized level, one can show that the backreaction always causes the throat to lengthen [35, 46], so that it cannot be traversed in any state of the decoupled system, as expected.

We will consider a deformation of the Hamiltonian that turns on at some time t_0 in (1.2).² At the linearized level, this has the same effect as changing the state to the future of t_0 . Now there is no reason for the above cancellation to occur since T_{UU} along $U < 0$ is unchanged. Therefore, one expects that generically by an appropriate choice of sign one will render the Einstein-Rosen bridge traversable, as long as the deformation couples the two boundaries.³

The simplest option in the large N limit is a double trace deformation. This has the effect of modifying the boundary conditions for the dual scalar field, such that some amplitude of a wave hitting one boundary will be transmitted to the opposite one. This does not change the eternal black hole solution classically, but results in a quantum correction to the matter stress tensor.

In order to be sure that the configuration is an allowed one, we choose the deformation to be relevant. Then it will be a renormalizable deformation of the CFT, and the dual

²We do not consider the case of a time-independent interaction, in order to prevent the quantum state from becoming non-regular on the past horizon.

³A deformation of only H_R has the same effect on the ANE as a change in the state, by bulk causality, since the past causal cone of the deformation does not intersect the $V = 0$ null sheet. This again agrees with the fact that when the boundaries are decoupled, no traversable wormhole can exist.

geometry will not be modified by backreaction in an uncontrolled way at the AdS boundaries. Also, heuristically, the effect of such a deformation coupling the two CFT's should be strong in the IR, which suggests that it renders the deep interior traversable.

Recall that the conformal weight of a scalar operator \mathcal{O}_i is given by $\Delta = \frac{d}{2} \pm \sqrt{\left(\frac{d}{2}\right)^2 + M^2}$, where M is the mass of the bulk field, and the plus or minus sign depends on the choice of asymptotic boundary conditions. In the case $M^2 > 0$, only the plus sign leads to normalizable modes. However, unitarity in AdS space [12] allows a slightly tachyonic bulk field: $M^2 > -\left(\frac{d}{2}\right)^2$, in which modes of both signs are normalizable and we are free to choose either one. To have a relevant deformation, we start with the alternative boundary condition, associated with the minus sign.

A brief overview of this paper is as follows. In section 2, we calculate the bulk two-point function with the modified Hamiltonian at linear order in h . In section 3, we use the point-splitting method to calculate T_{UU} on the $V = 0$ slice. Numerical result shows that T_{UU} is rendered negative by our boundary interaction. We find an analytic expression for $\int dUT_{UU}$, which is negative for all $0 < \Delta < 1$. In section 4 we calculate the energy and entropy of the resulting CFT state, and describe their holographic bulk duals. In section 5, we discuss the properties of this traversable wormhole and propose a quantum teleportation interpretation in the ER=EPR context. The appendix is a detailed calculation of $\int dUT_{UU}$.

Throughout we use units where $c = \hbar = 1$.

2 Modified bulk two-point function

For simplicity, we consider the eternal BTZ black hole [4, 5] (for a review, see [15]), whose metric is

$$ds^2 = -\frac{r^2 - r_h^2}{\ell^2} dt^2 + \frac{\ell^2}{r^2 - r_h^2} dr^2 + r^2 d\phi^2 \quad (2.1)$$

The inverse temperature of the BTZ black hole is determined by its horizon radius r_h as $\beta = 2\pi\ell^2/r_h$. Here and below we set AdS length ℓ to 1. Without any deformation of the Hamiltonian, the bulk free field two-point function in the BTZ background with $r^{-\Delta}$ fall-off was first derived by the mode sum method in [29].

In right wedge, it is

$$\langle \varphi_R(x) \varphi_R(x') \rangle_0 = \frac{1}{2^{3-\Delta} \pi} (G_+ + G_-) (G_+^{-1} + G_-^{-1})^{1-2\Delta} \quad (2.2)$$

where

$$G_{\pm} \equiv \left(\frac{rr'}{r_h^2} \cosh r_h \Delta \phi \pm 1 - \frac{(r^2 - r_h^2)^{1/2} (r'^2 - r_h^2)^{1/2}}{r_h^2} \cosh r_h \Delta t \right)^{-1/2}. \quad (2.3)$$

The bulk field operator $\varphi_R(x)$ in the eternal black hole background can be understood as a non-local CFT operator [41]. In particular, $\varphi_R(x)$ can be expanded in terms of the right boundary dual operator as

$$\varphi_R(t, r, \phi) = \int_{\omega > 0} d\omega dm \left(f_{\omega m}(r) e^{-i\omega t + im\phi} \mathcal{O}_{\omega m} + f_{\omega m}^*(r) e^{i\omega t - im\phi} \mathcal{O}_{\omega m}^\dagger \right) \quad (2.4)$$

where $f_{\omega m}(r)e^{-i\omega t+im\phi}$ are bulk positive frequency normalizable modes approaching $r^{-\Delta}$ when $r \rightarrow \infty$ and $\mathcal{O}_{\omega m}$ is the boundary annihilation operator defined by

$$\mathcal{O}(t, \phi) = \int d\omega dm \left(e^{-i\omega t+im\phi} \mathcal{O}_{\omega m} + e^{i\omega t-im\phi} \mathcal{O}_{\omega m}^\dagger \right). \quad (2.5)$$

Therefore, the bulk to boundary correlation function is given by

$$\begin{aligned} K_\Delta(r, t, \phi) &\equiv \langle \varphi_R(t, r, \phi) \mathcal{O}(0, 0) \rangle = \lim_{r' \rightarrow \infty} r'^{\Delta} \langle \varphi_R(t, r, \phi) \varphi_R(0, r', 0) \rangle_0 \\ &= \frac{r_h^\Delta}{2^{\Delta+1} \pi} \left(-\frac{(r^2 - r_h^2)^{1/2}}{r_h} \cosh r_h t + \frac{r}{r_h} \cosh r_h \phi \right)^{-\Delta}, \end{aligned} \quad (2.6)$$

where we used translation symmetry in t and ϕ to move (t', ϕ') to the boundary origin. This expression is real only when (r, t, ϕ) is space-like separated from the boundary origin. For time-like separation, general analytic properties of Wightman functions imply that one should change t to $t - i\epsilon$, which assigns a phase of $e^{-i\pi\Delta}$ when $t > 0$ and of $e^{i\pi\Delta}$ when $t < 0$.

Now we consider the time dependent modified Hamiltonian of (1.2):

$$\delta H(t) = - \int d\phi h(t, \phi) \mathcal{O}_R(t, \phi) \mathcal{O}_L(-t, \phi), \quad (2.7)$$

where $h(t, \phi) = 0$ when $t < t_0$. Using evolution operator $U(t, t_0) = \mathcal{T} e^{-i \int_{t_0}^t dt \delta H(t)}$ in interaction picture, the bulk two-point function is

$$\langle \varphi_R^H(t, r, \phi) \varphi_R^H(t', r', \phi') \rangle = \langle U^{-1}(t, t_0) \varphi_R^I(t, r, \phi) U(t, t_0) U^{-1}(t', t_0) \varphi_R^I(t', r, \phi) U(t', t_0) \rangle \quad (2.8)$$

where superscripts H and I represent Heisenberg and interaction picture respectively. To leading order in \hbar , (2.8) is (suppressing r and ϕ coordinates and omitting I)

$$\begin{aligned} G_h &\equiv -i \int_{t_0}^t dt_1 h(t_1) \langle [\mathcal{O}_L(-t_1) \mathcal{O}_R(t_1), \varphi_R(t)] \varphi_R(t') \rangle \\ &\quad - i \int_{t_0}^{t'} dt_1 h(t_1) \langle \varphi_R(t) [\mathcal{O}_L(-t_1) \mathcal{O}_R(t_1), \varphi_R(t')] \rangle \\ &\simeq -i \int_{t_0}^t dt_1 h(t_1) \langle \varphi_R(t') \mathcal{O}_L(-t_1) \rangle \langle [\mathcal{O}_R(t_1), \varphi_R(t)] \rangle + (t \leftrightarrow t') \\ &= i \int_{t_0}^t dt_1 h(t_1) K_\Delta(t' + t_1 - i\beta/2) [K_\Delta(t - t_1 - i\epsilon) - K_\Delta(t - t_1 + i\epsilon)] + (t \leftrightarrow t') \\ &= 2 \sin \pi \Delta \int dt_1 h(t_1) K_\Delta(t' + t_1 - i\beta/2) K_\Delta^r(t - t_1) + (t \leftrightarrow t') \end{aligned} \quad (2.9)$$

where in the second line we used large N factorization and causality, in that \mathcal{O}_L commutes with any φ_R , in the third line we used the KMS condition [23]

$$\langle \mathcal{O}_R(t) \mathcal{O}_L(t') \rangle_{tfd} = \langle \mathcal{O}_R(t) \mathcal{O}_R(t' + i\beta/2) \rangle_{tfd} \quad (2.10)$$

and in the last line K_Δ^r is the retarded correlation function

$$K_\Delta^r(t, r, \phi) = |K_\Delta(t, r, \phi)| \theta(t) \theta \left(\frac{(r^2 - r_h^2)^{1/2}}{r_h} \cosh r_h t - \frac{r}{r_h} \cosh r_h \phi \right) \quad (2.11)$$

One can also derive (2.9) using the bulk mode sum method with modified boundary conditions. This approach would allow one to compute the stress tensor for finite h , not just perturbatively. According to Lorentzian AdS/CFT, the double trace deformation [8, 58], from the point of view of the right wedge, is equivalent to a source term $h(t, \phi)\mathcal{O}_L(-t, \phi)$, for $\mathcal{O}_R(t)$, activating the initially frozen fall-off component of the bulk field. The same applies to the left wedge. Therefore the asymptotic behavior of a global bulk mode φ living in the entire eternal black hole should satisfy

$$\varphi(r \rightarrow \infty_R) \rightarrow \alpha_R(t, \phi)r^{-\Delta} + \beta_R(t, \phi)r^{-2+\Delta}, \quad \beta_L(t, \phi) = h(-t, \phi)\alpha_R(-t, \phi) \quad (2.12)$$

$$\varphi(r \rightarrow \infty_L) \rightarrow \alpha_L(t, \phi)r^{-\Delta} + \beta_L(t, \phi)r^{-2+\Delta}, \quad \beta_R(t, \phi) = h(t, \phi)\alpha_L(-t, \phi) \quad (2.13)$$

where the subscript 1 is for right wedge and 2 is for left wedge.

The thermofield double state of the eternal black hole is the vacuum state in the Kruskal patch [30]. This is analogous to the relation between the Minkowski vacuum and the Rindler thermofield double state [51]. Choosing the appropriate global bulk modes $H_{\omega m}^{(\pm)}$,⁴ and applying the method of [30], we can construct φ as

$$\varphi(x) = \int_{\omega>0} d\omega dm (H_{\omega m}^{(+)}(x)b_{\omega m}^{(+)} + H_{\omega m}^{(-)}(x)b_{\omega m}^{(-)\dagger} + h.c.) \quad (2.14)$$

where $b_{\omega m}^{(\pm)}$ are annihilation operators used to define the vacuum. We find the two-point function in this vacuum is the same as (2.9) up to normalization. Since the calculation is quite involved, we do not include it in this paper.

3 1-loop stress tensor

The stress tensor is given by variation of action with respect to $g^{\mu\nu}$,

$$T_{\mu\nu} = \partial_\mu\varphi\partial_\nu\varphi - \frac{1}{2}g_{\mu\nu}g^{\rho\sigma}\partial_\rho\varphi\partial_\sigma\varphi - \frac{1}{2}g_{\mu\nu}M^2\varphi^2 \quad (3.1)$$

The 1-loop expectation value can be calculated by point splitting,

$$\langle T_{\mu\nu} \rangle = \lim_{x \rightarrow x'} \partial_\mu\partial'_\nu G(x, x') - \frac{1}{2}g_{\mu\nu}g^{\rho\sigma}\partial_\rho\partial'_\sigma G(x, x') - \frac{1}{2}g_{\mu\nu}M^2 G(x, x') \quad (3.2)$$

where $G(x, x')$ is 2-point function. In this formula, one must renormalize the stress tensor by subtracting the coincident point singularities from the 2-point function, which are given by the Hadamard conditions [42]. Since these are determined by the short distance dynamics, this subtraction is unchanged when we modify the boundary conditions, and it has no effect on the order h correction that we are interested in.

At leading order, as we reviewed in the Introduction, $\int dU T_{UU}$ is zero on the horizon $V = 0$. Indeed, the leading order two point function in the BTZ black hole is (2.2) where ϕ has periodicity 2π and all $\Delta\phi + 2\pi n$ images are summed. The only coincident point pole

⁴This step is very tricky because at order h , the $r^{-\Delta}$ component is not constrained by the deformation. The only requirement is that the modified two point function must be regular on horizon. We were able to find a choice to reproduce (2.9) up to normalization.

comes from the $n = 0$ component. Summing over the other n components, one finds that in Kruskal coordinates the leading order stress tensor $T_{UU} \sim O(V^2)$ in the $V \rightarrow 0$ limit, so that $T_{UU} = 0$ along the horizon.

The subleading 2-point function is given by (2.9). Note that $h(t, \phi)$ is dimensionful and its dimension is $2 - 2\Delta$ because in (2.7) \mathcal{O} has dimension Δ .⁵ Moreover, since $h(t, \phi)$ is a boundary CFT smearing function, it should not depend on any bulk length scale (e.g. r_h and ℓ) explicitly but only on the inverse temperature β . Let us assume that $h(t, \phi)$ is uniform over ϕ :

$$h(t, \phi) = \begin{cases} h(2\pi/\beta)^{2-2\Delta} & t \geq t_0 \\ 0 & t < t_0 \end{cases} \quad (3.3)$$

where h is a dimensionless constant. In Kruskal coordinates

$$e^{2r_h t} = -\frac{U}{V}, \quad \frac{r}{r_h} = \frac{1 - UV}{1 + UV} \quad (3.4)$$

the change in the 2-point function is

$$G_h = C_0 \left(\frac{2\pi}{\beta}\right)^{2\Delta-2} r_h \int \frac{dU_1}{U_1} d\phi_1 h(U_1, \phi_1) \left(\frac{1 + UV}{U/U_1 - VU_1 - (1 - UV) \cosh r_h(\phi - \phi_1)} \right)^\Delta \\ \times \left(\frac{1 + U'V'}{U'U_1 - V'/U_1 + (1 - U'V') \cosh r_h(\phi' - \phi_1)} \right)^\Delta + (U, V, \phi \leftrightarrow U'V', \phi') \quad (3.5)$$

where $C_0 = \frac{r_h^{2\Delta-2} \sin \Delta\pi}{2(2\Delta\pi)^2} \left(\frac{2\pi}{\beta}\right)^{2-2\Delta}$ and we transformed the integral over t_1 to Kruskal coordinates in which the boundary is $U_1 V_1 = -1$. Note that this result applies to both the black hole and black brane cases because the integration of ϕ_1 over 0 to 2π and summing over n with modification $\phi_1 \rightarrow \phi + 2\pi n$ is equivalent to the integration of ϕ_1 over the whole real axis. Since we only focus on T_{UU} component on the horizon $V = 0$ and the derivative on U and U' in (3.2) has nothing to do with the value of V and V' , we can take both points to the horizon first, namely $V = V' = 0$. Similarly, we can take $\phi = \phi'$ first for simplicity. Since $h(t_1, \phi_1)$ is uniform in ϕ_1 , ∂_ϕ is still a Killing vector of the system and therefore G_h should not depend on ϕ . Defining $y = \cosh r_h(\phi_1 - \phi)$, on horizon we have

$$G_h = hC_0 \int_{U_0}^U \frac{dU_1}{U_1} \int_1^{\frac{U}{U_1}} \frac{2dy}{\sqrt{y^2 - 1}} \left(\frac{U_1}{U - U_1 y}\right)^\Delta \left(\frac{1}{U'U_1 + y}\right)^\Delta + (U \leftrightarrow U') \\ \equiv F(U, U') + F(U', U) \quad (3.6)$$

where $U_0 = e^{r_h t_0}$. The integral range of (3.6) is given by the step function in (2.11), which ensures that $U - U_1 y \geq 0$. Note that the integral in (3.6) is dimensionless. Since G_h has dimension 1 (φ_R has dimension $\frac{1}{2}$ in 3-dimension spacetime), if we restore ℓ in (3.6), we find the total length scale dependence of G_h is ℓ^{-1} .

Note that $g_{UU} = 0$ in the original BTZ geometry. By (3.2), T_{UU} on horizon is

$$T_{UU} = \lim_{U' \rightarrow U} \partial_U \partial_{U'} (F(U, U') + F(U', U)) = 2 \lim_{U' \rightarrow U} \partial_U \partial_{U'} F(U, U') \quad (3.7)$$

⁵Here we implicitly defined the unit length angular coordinate $x \equiv \phi\ell$. Taking the limit $r \rightarrow \infty$ in BTZ metric (2.1), the boundary metric is flat $ds_b^2 = -dt^2 + dx^2$.

where we should note the dimension of T_{UU} is the same as G_h because U is dimensionless. Since the integration ranges are only functions of U , we can take the U' derivative before evaluating the integral

$$T_{UU} = -4h\Delta C_0 \lim_{U' \rightarrow U} \partial_U \int_{U_0}^U dU_1 \int_1^{\frac{U}{U_1}} \frac{dy}{\sqrt{y^2 - 1}} \frac{U_1^\Delta}{(U - U_1 y)^\Delta (U' U_1 + y)^{\Delta+1}} \quad (3.8)$$

Defining a new variable $x = \frac{y-1}{U/U_1-1}$ and integrating over x we get

$$T_{UU} = -\frac{4h\Delta C_0 \Gamma(\frac{1}{2}) \Gamma(1-\Delta)}{\sqrt{2} \Gamma(\frac{3}{2}-\Delta)} \lim_{U' \rightarrow U} \partial_U \int_{U_0}^U dU_1 \frac{F_1\left(\frac{1}{2}; \frac{1}{2}, \Delta+1; \frac{3}{2}-\Delta; \frac{U_1-U}{2U_1}, \frac{U_1-U}{U_1(1+U'U_1)}\right)}{U_1^{-\Delta+1/2} (U-U_1)^{\Delta-1/2} (1+U'U_1)^{\Delta+1}} \quad (3.9)$$

where we used the integral representation of Appell hypergeometric function. The integral over U_1 is finite as long as $\Delta - 1/2 < 1$, namely $\Delta < 3/2$, because in the integrated region, the only potentially divergent point is around $U_1 \rightarrow U$ from below since F_1 is a complete function when $\Delta < 3/2$. In particular, when $U_1 \sim U$, $F_1 \rightarrow 1$, which implies $\Delta < 3/2$ is the sufficient and necessary condition for integrability. Defining a new variable $z = \frac{U_1-U_0}{U-U_0}$, the domain of integration in (3.9) becomes 0 to 1 and therefore we can exchange the order of ∂_U and $\int dz$. After differentiating w.r.t. U , and restoring the variable U_1 , we get

$$T_{UU} = -\frac{2h\Delta C_0 \Gamma(\frac{1}{2}) \Gamma(1-\Delta)}{\Gamma(\frac{3}{2}-\Delta)} \int_{U_0}^U \frac{dU_1 U_1^{2\Delta} (f_1 + f_2 + f_3)}{(U-U_0)(U-U_1)^{\Delta-1/2} (1+U_1^2)^{\Delta+1} U^{\Delta+1} (U+U_1)^{1/2}} \quad (3.10)$$

where

$$f_1 = \frac{-2\Delta(UU_1^2 + U_0) + 3UU_0U_1 + U_0 + 2U_1}{1 + UU_1} F_1\left(1-\Delta, \frac{1}{2}, 1+\Delta, \frac{3}{2}-\Delta, u, v\right) \quad (3.11)$$

$$f_2 = \frac{2(1+\Delta)(U-U_1)(U_0 + 2UU_0U_1 - UU_1^2)}{(2\Delta-3)U(1+U_1^2)(1+UU_1)} F_1\left(1-\Delta, \frac{1}{2}, 2+\Delta, \frac{5}{2}-\Delta, u, v\right) \quad (3.12)$$

$$f_3 = \frac{U_0(U-U_1)}{(2\Delta-3)(U+U_1)} F_1\left(1-\Delta, \frac{3}{2}, 1+\Delta, \frac{5}{2}-\Delta, u, v\right) \quad (3.13)$$

$$u = \frac{U-U_1}{U+U_1}, \quad v = \frac{U-U_1}{U(1+U_1^2)} \quad (3.14)$$

Performing the final integral numerically, we plot the result in figure 2a.

In the figure, we see that the null energy is negative after we turn on the insertion at $U_0 = 1$ if we take positive h . Physically, this means the light-like ray $V = 0$ becomes time-like after U_0 and a spaceship that enters early enough may escape the black hole!

One may note that when $\Delta < 1/2$, T_{UU} is finite but when $\Delta > 1/2$, T_{UU} is singular near insertion time U_0 . However, this singularity is not essential because it is integrable, as we will see later when we calculate $\int dUT_{UU}$ along the horizon $V = 0$. Indeed, the classical solution of Einstein equations for a shockwave insertion in the bulk in Kruskal coordinates contains a delta function, which is also an integrable singularity [46]. One

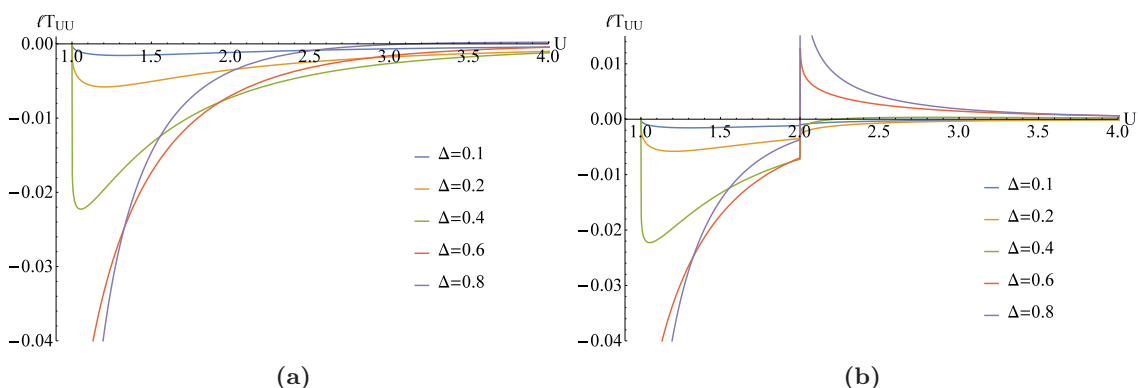


Figure 2. (a) shows the null energy along the horizon when the interaction is turned on at $U = U_0 = 1$ and never shut off, with our choice for the sign of the coupling h ; (b) shows the case where it is turned on at $U = U_0 = 1$ and turned off at $U = U_f = 2$. In both cases, $h = 1$. We see clearly in both (a) and (b) that T_{UU} becomes negative after turn-on; in (b) T_{UU} becomes positive after turn-off. Blue is for $\Delta = 0.1$; yellow is for $\Delta = 0.2$; green is for $\Delta = 0.4$; pink is for $\Delta = 0.6$; purple is for $\Delta = 0.8$.

might also worry that the derivative of g_{UU} and the Riemann curvature are singular at the turn-on and turn-off times, although T_{UU} and $\int dU T_{UU}$ are not. However, this is simply due to the fact that we turned the insertion on and off as a step function. If this process were taken to be smooth enough, there would be no singularity.

To see the late time behavior, we can use the z variable to rewrite (3.10) in the large U limit. In this limit, f_1 dominates among all f_i 's in (3.10). Using the identity $F_1(a; b, b'; c; z, 0) = {}_2F_1(a, b; c; z)$ we obtain

$$T_{UU} \rightarrow \frac{4h\Delta^2 C_0 \Gamma(\frac{1}{2}) \Gamma(1-\Delta)}{\Gamma(\frac{3}{2}-\Delta) U^{2\Delta+2}} \int_0^1 \frac{dz z^{2\Delta+1} {}_2F_1\left(1-\Delta, \frac{1}{2}, \frac{3}{2}-\Delta, \frac{1-z}{1+z}\right)}{((z+\epsilon)^2 + \epsilon)^{\Delta+1} (1-z)^{\Delta-1/2} (1+z)^{1/2}} \rightarrow 0_+ \quad (3.15)$$

where ϵ is a small number of order U^{-1} and which implies that T_{UU} becomes positive and decays to zero at late times.

If we turn off the interaction at some finite time U_f , when $U > U_f$, we can safely pass ∂_U into the U_1 integral, which leads to

$$T_{UU} = -\frac{4h\Delta C_0 \Gamma(\frac{1}{2}) \Gamma(1-\Delta)}{\Gamma(\frac{1}{2}-\Delta)} \int_{U_0}^{U_f} dU_1 \frac{U_1^{2\Delta+1} F_1\left(-\Delta; \frac{1}{2}, \Delta+1; \frac{1}{2}-\Delta; \frac{U-U_1}{U+U_1}, \frac{U-U_1}{U+U_1^2}\right)}{(U-U_1)^{\Delta+1/2} (U+U_1)^{1/2} U^{\Delta+1} (1+U_1^2)^{\Delta+1}} \quad (3.16)$$

In deriving (3.16), we used a property of the derivative of the Appell hypergeometric function and equation (7a) in [45]. The numerical result is plotted in figure 2b.

In this figure, we see that after turning off the interaction, T_{UU} has a jump and becomes positive at late times. In particular, when $\Delta > 1/2$, T_{UU} becomes divergent again right after U_f . Fortunately, it is again an integrable divergence which should not cause any

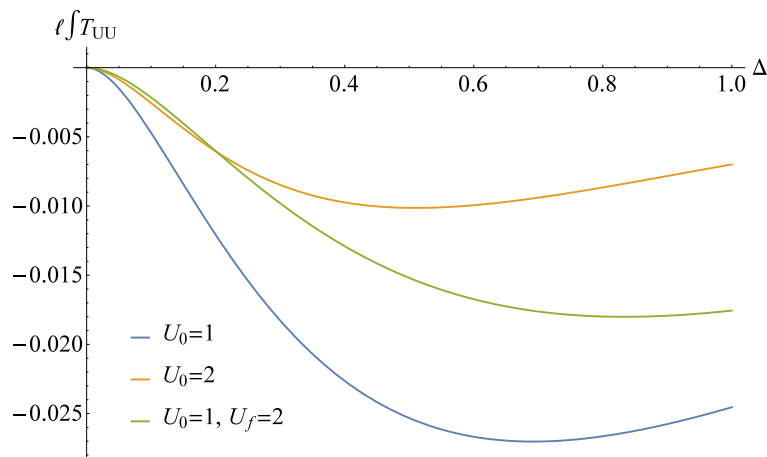


Figure 3. $\int dUT_{UU}$ as a function of Δ ; blue is for $U_0 = 1$; yellow is for $U_0 = 2$; green is for $U_0 = 1$ and $U_f = 2$.

physical problem. By the identity [43]:

$$F_1(a; b, b'; c; x, y) = \sum_{m \geq 0} \frac{(a)_m (b)_m}{m! (c)_m} x^m {}_2F_1(a + m, b'; c + m; y) \quad (3.17)$$

the late time behavior can be analyzed by taking the $U \rightarrow \infty$ limit in (3.16):

$$T_{UU} \sim \frac{4h\Delta^2 C_0}{U^{2\Delta+2}} \log U \log \frac{U_f}{U_0} \rightarrow 0_+$$

Again, we find T_{UU} becomes positive after some time and decays to zero. In both late time analyses, T_{UU} decays like $U^{-2\Delta-2}$, which validates the assumption that Uh_{UU} and $\partial_U h_{\phi\phi}$ vanish when $U \rightarrow \infty$ in (1.3).

In the above discussion, we see that at some finite time T_{UU} becomes positive whether or not we turn off the insertion, which might appear dangerous for the fate of the worm hole. The crucial diagnostic is the sign of the integral of T_{UU} over the whole $V = 0$ slice. This is what determines whether a light ray on horizon eventually reaches the boundary at spatial infinity.

It looks horrible to integrate U in (3.10) from U_0 to infinity. Interestingly and surprisingly, by some tricks, we can get a closed form for it (see appendix A):

$$\int_{U_0}^{\infty} dUT_{UU} = -\frac{h\Gamma(2\Delta + 1)^2}{2^{4\Delta}(2\Delta + 1)\Gamma(\Delta)^2\Gamma(\Delta + 1)^2\ell} \frac{{}_2F_1\left(\frac{1}{2} + \Delta, \frac{1}{2} - \Delta; \frac{3}{2} + \Delta; \frac{1}{1+U_0^2}\right)}{(1 + U_0^2)^{\Delta+1/2}} \quad (3.18)$$

If we turn off the interaction at U_f , the integral is just the difference between $\int_{U_0}^{\infty} dUT_{UU}$ and $\int_{U_f}^{\infty} dUT_{UU}$. We plot the result as a function of Δ in figure 3.

In this figure, we see that for all Δ values from 0 to 1, the integral of T_{UU} is always negative, which demonstrates the existence of a traversable wormhole. Furthermore, the earlier we turn on the insertion, the larger the effect is. In particular, even if T_{UU} becomes positive in late times, the wormhole still exists since the integral of T_{UU} remains negative.

Note that $\Delta = 0$ is a special case where $\int dU T_{UU} = 0$. Indeed, the only $\Delta = 0$ operator in CFT is the identity and of course adding the product of identity operators to Hamiltonian has no effect on the system.

4 Holographic energy and entropy

In this section we will consider the implications of a traversable wormhole for the holographic entanglement entropy conjecture, which in this context relates the entanglement entropy between the two boundary CFT's to the area/entropy of certain extremal surfaces in the bulk theory [7, 18, 20, 28, 44].

As a preliminary, we discuss the change of energy of the CFT state. Long after the interaction is shut off, the system returns to thermal equilibrium. Thus the final horizon area can be determined from the energy of the system, measured on the left or the right. It is straightforward to check that, in our state, the energy decreases at linear order in h with the sign choice that rendered the wormhole traversable:

After deforming the Hamiltonian ($t > t_0$), the state in Schrödinger picture is

$$|\Psi(t)\rangle = e^{-iH_0(t-t_0)}U(t, t_0) |\text{tdf}\rangle. \tag{4.1}$$

Expanding $U(t, t_0)$ to leading order in $h(t)$ given by (3.3), we find that the change in the energy on the right is

$$\begin{aligned} \delta E_R &= i \int_{t_0}^t dt_1 h(t_1) \langle \text{tdf} | [\delta H(t_1), H_R] | \text{tdf} \rangle \\ &= \int_{t_0}^t dt_1 d\phi h(t_1) \langle \text{tdf} | \partial_t \mathcal{O}_R(t_1, \phi) \mathcal{O}_L(-t_1, \phi) | \text{tdf} \rangle \\ &= \frac{hr_h^2}{2^{\Delta+1}\ell^3} \sum_n \left(\frac{1}{(\cosh 2r_h t + \cosh 2\pi r_h n)^\Delta} - \frac{1}{(\cosh 2r_h t_0 + \cosh 2\pi r_h n)^\Delta} \right) \end{aligned} \tag{4.2}$$

where in the second line we used the Heisenberg equation and in last line the boundary two-point function is obtained by taking limit $r \rightarrow \infty$ in (2.2) where ϕ has period 2π , and all of its images are summed over in the global BTZ black hole.⁶ If the interaction shuts off at t_f , the energy obviously becomes constant for $t > t_f$, and t in (4.2) is replaced by t_f . Therefore, the effect of the interaction with $h > 0$ is to reduce the energy. Note that if there are any UV divergences in the energy they cannot appear at linear order in h , since the interaction involves just one field in each CFT.

At least at first order in h , the entropy of entanglement S_{EE} between the left and right boundaries should also be well-defined (and time dependent) even during the period of time when the interaction is turned on, if one thinks of the state as evolving by the deformed Hamiltonian in the original tensor product Hilbert space. By the first law of entanglement, at linear order in h , the change in S_{EE} is equal to $\beta\delta H_R$, thus it also decreases until the turn-off time t_f after which it remains constant (as it must under decoupled unitary evolution on the left and right).

⁶We consider global AdS here so that the total energy is finite.

The change in S_{EE} is $\mathcal{O}(1)$ in a $1/c \sim 1/N \sim G_N/L_{\text{AdS}}$ expansion.⁷ At this order, in the bulk interpretation S_{EE} has two parts, namely the small gravitational correction to the area $A/4G_N$ of the extremal surface, and the entanglement entropy of bulk fields S_{bulk} on the spacelike slice from the extremal surface to the boundary slice at time t [7, 20].

In our situation, causality implies that the geometry near the bifurcation surface is unaffected by the perturbation. Thus, at order h , the area of the quantum extremal surface is unchanged from the original state. On the other hand, S_{bulk} has nonlocal aspects. Therefore, the decrease of S_{EE} at first order must be entirely due to a corresponding decrease in S_{bulk} evaluated at the bifurcation surface.

However, after the time t_{trav} on the boundary, the bifurcation surface (E_1 of figure 4) is no longer spacelike to the AdS boundary, and one cannot define a bulk entanglement wedge using the surface E_1 . This would render S_{bulk} ill-defined. The resolution is to use the *quantum extremal* surface.

More generally, it was proposed in [18] (and proven in [17]) that, at general orders in $1/N$, one should consider the entropy outside the quantum extremal surface, obtained by extremizing the total generalized entropy $S_{\text{gen}} = A/4G_N + S_{\text{bulk}}$. When calculating the $\mathcal{O}(1)$ piece of the entropy, these two prescriptions agree on the value of the entropy but [16, 18] argued that the location of the quantum extremal surface is also physically important, because it provides a natural boundary for how much of the bulk can be reconstructed from the CFT state on a single boundary. One useful constraint on the location quantum extremal surface is the GSL, which states that S_{gen} is nondecreasing on any future horizon.

On a Cauchy slice prior to the time when the interaction is turned on, the geometry and bulk quantum state are that of the Hartle-Hawking state. Thus the quantum (and classical) extremal surface is located at the bifurcation surface of the original black hole (E_1 of figure 4). On the other hand, after the interaction is over, the bulk quantum state of the fields changes and thus the quantum extremal surface must move. By left-right symmetry of the spacetime (together with the fact that the joint state of the entire system is pure so that S_{EE} is the same on both sides) it can only move along the vertical axis of symmetry of the spacetime. Also, the GSL implies that the new location must be on or behind the causal horizon [18], because otherwise it lies on a future horizon whose S_{gen} is generically increasing.

In fact, at first order in h , the GSL implies that the quantum extremal surface must lie exactly at the point E_2 in figure 4, where the two future horizons intersect. For since the GSL is true in every state [56], and saturated for the Hartle-Hawking state, it must also be saturated for any first order perturbation to the Hartle-Hawking state [54]. But if S_{gen} is stationary along two linearly independent normal directions of E_2 , then it must be a quantum extremal surface.

As noted above, the area and bulk entanglement entropies are identical for E_1 and E_2 at linear order in h (as long as they are well-defined). Any effects arising from differences between E_1 and E_2 are suppressed by additional powers of h .

⁷These are the correct scaling relations for 2+1 dimensional bulk. In other dimensions, the scaling with G_N is the same, but the scaling with the number of species N may vary.

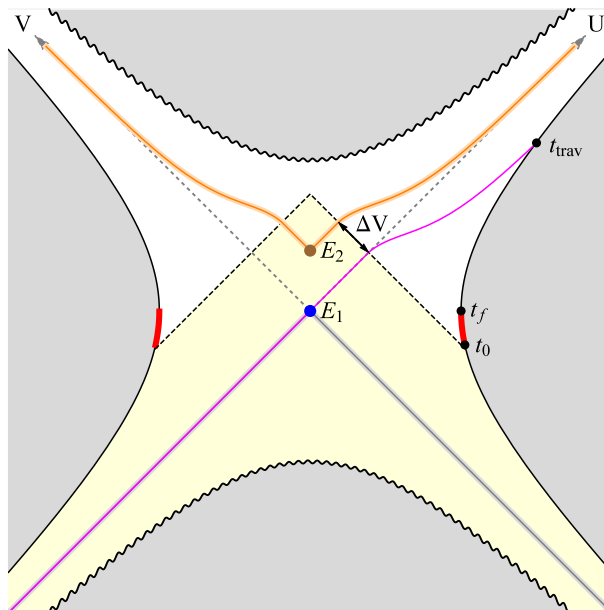


Figure 4. The throat size is $\Delta V \sim h$. The red thick interval on the boundary is the duration of the deformation beginning at t_0 and ending at t_f . The metric in the light yellow region is unchanged and only that of the white region will have a nonzero backreaction correction. The orange thick curve is the future event horizon and the grey thick curve is the past event horizon. E_1 is the original bifurcation surface. E_2 is the location where the right and left future horizons cross. The magenta curve is a null ray that passes through wormhole, deviating to right boundary.

At second order in h , the GSL should no longer be saturated on the future horizon. Hence S_{bulk} is increasing with time at E_2 , and the quantum extremal surface will instead be located slightly above the point E_2 .

We have not followed the evolution of the quantum extremal surface at intermediate times, but it seems that it must gradually move upwards from E_1 to its final location above E_2 . After the interaction is over the boundary evolution is unitary, and hence neither S_{EE} nor the quantum extremal surface changes.

Ref. [18] argued that the quantum extremal surface should always be spacelike to its corresponding CFT region. In a sense this continues to be true, since E_1 is spacelike to all the boundary points prior to turning on the interaction, while E_2 is spacelike to all the points after the interaction is turned off. But neither one is spacelike to the entire boundary for all time. For example, a unitary operator applied to the right boundary at sufficiently early times might affect the value of $S_{\text{gen}}(E_2)$, and hence the right CFT entropy after the interaction. But that does not contradict any of the properties of the right CFT, since it does not have unitary time evolution (independent of the left CFT) during the period of the interaction.

Note that, if we assume that our holographic entropy prescription is correct when the CFT's are not coupled, it must necessarily also be correct when the CFT's are coupled. Before the interaction is turned on, we can simply consider the Hartle-Hawking spacetime as if there were no interaction. Similarly, after the interaction is over, we can consider a

new spacetime which is dual to extrapolating the final state backwards in time, without any interaction. Neither of these spacetimes corresponds to a traversable wormhole, but they can be used for purposes of calculating S_{EE} before or after the interaction is turned on. It is only when these two spacetimes are patched together, that they are seen to be a traversable wormhole geometry.

5 Discussion

We have demonstrated that the Einstein-Rosen bridge of a BTZ black hole becomes slightly traversable after the addition of a two-boundary coupling. (We expect that a similar effect also occurs in $D > 3$ bulk spacetime dimensions, although it is harder to calculate the exact form of the stress-tensor.)

From (3.18), we see that the integral $\int dUT_{UU}$, giving the deviation of null rays from the horizon, is proportional to h , which implies that the wormhole opens up by an amount (in units where $\hbar = 1$)

$$\Delta V \sim \frac{hG_N}{R^{D-2}} \tag{5.1}$$

where ΔV is the difference of V coordinate between the future horizon and the first lightray which can get through the wormhole (see figure 4), and we assume that the black hole radius r_h , the AdS length L_{AdS} , and the amount of time Δt the interaction is turned on for are all of the same characteristic length scale R .

The wormhole is only open for a small proper time in the interior region. This is quite different from the usual static wormhole solutions which do not have event horizons (e.g. [38]). Nevertheless, radial lightrays originating on the boundary at arbitrarily early times will cross through the portal to the other side; in this sense the wormhole is open at arbitrarily early boundary times on either side.

A (test) astronaut from one boundary can only go through the wormhole before it closes, and she reaches the other boundary long after the boundary-boundary interaction is turned on. One should note that since the coupling we add breaks the Killing symmetry $H_L - H_R$, there is no way to boost her back to a time before she entered the worm hole. Thus the way we glue the two boundaries fixes the relative time coordinate between them, excluding the possibility of having closed time-like curves [39]. Note that the traversable throat size depends on the strength of the coupling and a signal transmitted through the wormhole is only received at the other end after a very long time delay if the gravitational effects of the coupling are small. Furthermore, the thermofield double state that we require is an extremely fine tuned state, so it would be very difficult to prepare such a configuration in which the astronaut could enter at early times from the left.

We have not yet considered the backreaction on the geometry coming from an actual (non-test) astronaut traversing the wormhole throat. An object travelling at light speed from left to right contributes to T_{VV} but not to T_{UU} , so at the level of linearized gravity it will prevent objects from traversing in the *other direction* (i.e. from right to left) but it will have no tendency to close the wormhole in the same direction that it is travelling. This

suggests that the objects can still traverse the wormhole even after taking into account their own gravitational back-reaction.⁸

Another question concerns the interaction of the astronaut with the negative energy pulse of radiation travelling in the other direction. In the frame of reference defined by Kruskal coordinates, a quantum traversing the wormhole must be blueshifted up to a frequency $1/\Delta V$, while the pulse coming in the other direction has a frequency of order $1/R$. Here we are assuming that the interaction is turned on for about one light-crossing time R , and that there is no other time scale of relevance in the problem. Although an incoming pulse with negative total energy is not allowed in classical scattering problems, we will nevertheless attempt to build intuition by comparing the situation to a normal field theory scattering problem. The center-of-mass energy scale of the collision is given by

$$\sqrt{s} \sim \sqrt{\frac{R^{D-4}}{hG_N}}. \tag{5.2}$$

Since $G_N \sim L_{\text{planck}}^{D-2}$, the center-of-mass energy is below the Planck scale in $D = 3$ (i.e. a BTZ black hole with any extra dimensions compactified at the Planck scale) but not when $D > 3$. However, even in higher dimensions we do not expect that full quantum gravity effects will be important. We nevertheless expect that it is legitimate to use the eikonal approximation, in which one solves for the propagation of each particle on the background field generated by the other particle. This corresponds to resumming ladder Feynman diagrams, whose amplitude scales with various powers of

$$\frac{G_N s}{b^{D-4}} \sim h^{-1}, \tag{5.3}$$

where b is the impact parameter, and we have used the fact that $b \sim R$ (except for small tails of the wavefunction). Non-eikonal Feynman diagrams should be suppressed by additional powers of G_N relative to eikonal diagrams with the same s dependence [2]. Therefore we can consistently consider scenarios in which only the eikonal scattering is relevant, in which our calculation of the geometry shows that the wormhole is traversable if the particle is sent in before a time $\Delta t \approx R \ln(R/hL_{\text{planck}})$ prior to the interaction being turned on. However if the particle is sent in more than $\frac{3}{2}\Delta t$ prior to the interaction time, then the eikonal approximation breaks down and there may be large back-reaction, invalidating our analysis. When $h \sim 1$, these times are of order the scrambling time.

It is interesting to consider what would happen if the two black holes were in the same component of space, rather than in different asymptotic regions. If the black holes were in a suitably entangled state, they should be connected by an Einstein-Rosen bridge [35], with the QFT state near the horizon close to the Hartle-Hawking state. The direct boundary interaction could then be produced by propagation through the ambient spacetime — this

⁸Presumably there is *some* limit on how much information can get through, since the black hole on the other side cannot radiate more energy than its initial mass, but determining the precise limit would require going beyond the linearized regime. There might also be an interesting limit on the total amount of *information* which can get through the wormhole, coming from the Bousso bound [10] (see [21] for a proof in classical case) or its quantum generalization [11, 48].

would be the same as the interaction we studied, except with a time delay. A similar calculation would then lead to a traversable wormhole. The negative ANE could be understood as coming from the Casimir effect associated to the cycle in space going from one black hole to the other in the ambient space and then threading the wormhole. Of course, the effect would be enhanced if the signals sent between the black holes were directed and amplified (otherwise the Casimir energy would be extremely tiny if the black holes were far apart). No causal paradoxes would arise because the traversability depends on backreaction due to the existence of a casual path between the black holes in the ambient spacetime.

Since any infinite null geodesic which makes it through a wormhole must be chroral (as discussed in the Introduction), such wormholes do not enable one to travel faster than light over long distances through space. Hence traversable wormholes are like getting a bank loan: you can only get one if you are rich enough not to need it.

The traversable wormhole we found has an interesting interpretation in the context of ER=EPR [35]. Maldacena and Susskind conjectured that any pair of entangled quantum systems are connected by an Einstein-Rosen bridge (the non-traversable wormhole). The crucial difference in our work is that we allow interaction between the entangled systems, which is assumed to be negligible in ER=EPR. What we have shown is that in this case the Einstein-Rosen bridge can open to become a traversable wormhole.

Our example thus provides a way to operationally verify a salient feature of ER=EPR that observers from opposite sides of an entangled pair of systems may meet in the connected interior. Since in [35] any such meeting is trapped behind the horizon, it is not obvious how its occurrence could be confirmed by exterior or CFT measurements. What we found is that if, after the observers jump into their respective black holes, a boundary-boundary coupling is activated, then the Einstein-Rosen can be rendered traversable, and the meeting inside may be seen from the boundary. This seems to suggest that the ER=EPR wormhole connection was physically “real”. But since all measurements in the CFT description are governed by the rules of linear quantum mechanics, it seems like any explicit operational verification of the existence of the wormhole would also correspond to a linear quantum measurement. It might be interesting to check the compatibility of these ideas with the linearity of measurements made behind the horizon, discussed in [37].

What is the quantum information theory interpretation of such a traversable wormhole? A curious feature of the transmission of a qubit, Q , through the wormhole is that it appears to be sent “via the entanglement”, rather than directly by the inter-boundary coupling. (Note that the traversable portion of the wormhole is close to the bifurcation point, which describes the subspaces of the left and right Hilbert spaces that are the most entangled in the thermofield double state.) There are several ways to see that the quantum information of Q is not simply being sent directly through the boundaries. First, the commutator of Q (for example when it is first injected into the interior from the left boundary) with the interaction Hamiltonian is extremely small near the thermofield double state. Furthermore, at the time the interaction is activated, Q is in fact spacelike separated from the boundary in the bulk picture, so in the bulk approximation Q and \mathcal{O} are independent quantum variables. From the CFT perspective, this is because Q has thermalized into the left system before the $\mathcal{O}_L\mathcal{O}_R$ interaction is turned on, so no quantum information about Q appears to be

accessible to the operator \mathcal{O} . Of course, the boundary coupling is nevertheless crucial for the existence of the traversable wormhole.

This situation is somewhat analogous to what occurs in quantum teleportation. Entanglement alone cannot be used to transmit information, and no qubit, Q , from the left can traverse the bridge to the right if the left and right systems are dynamically decoupled. However, if additional classical information is sent from the left to the right, a qubit can be transmitted - this is referred to as quantum teleportation. Suppose Alice and Bob share a maximally entangled pair of qubits, A and B . Alice can then transmit the qubit Q to Bob by sending only the classical output of a measurement on the Q - A system. Depending on which of the 4 possible results are obtained, Bob will perform a given unitary operation on the qubit B , which is guaranteed to turn it into the state Q .

Unlike the usual description of quantum teleportation, in our example it is essential that the channel between the left CFT, A , and the right CFT, B , is a quantum one. For example, if one projected onto eigenstates of the operator \mathcal{O}_L , then the configuration would simply look like a particular quantum state (the projection of $|\text{tfd}\rangle$) evolving under the decoupled Hamiltonians together with an action by a purely right unitary, which can never lead to a traversable wormhole. This makes sense, because in the standard description of quantum teleportation, the measurement performed by Alice is a projection onto an eigenstate, which instantly results in the pattern of Q being contained in the system B . This would not be described by a physical motion through the wormhole in the bulk. Teleportation in this sense has been discussed in the dual gravity language by [37, 40, 49].

However, in the exact, fully quantum description of the quantum teleportation protocol, there is a particular dynamical process given by the unitary evolution $V = \sum_i P_i^{QA} U_i^B$ that governs the transmission of the “classical” information and the subsequent appropriate transformation of a qubit in the B quantum system. Here P_i^{QA} are a complete mutually exclusive set of projectors on the Q - A system that describe Alice’s measurement, and U_i^B is the unitary transformation performed by Bob given the data i . The classical information transmitted from Alice to Bob was encoded by the index i .

Treating V as a time dependent interaction Hamiltonian can result in negative ANE along the horizon if the original entanglement between A and B was well described by a large Einstein-Rosen bridge, which will render the wormhole traversable. This is a description in which the time scales and processes of decoherence and measurement by Alice are resolved, and treated as physical dynamical evolution. In such a “microscopic” description of quantum teleportation, the qubit Q must physically evolve from the left to the right. Of course in the limit that Alice’s measurement is essentially instantaneous and classical, the traversable window will be very small (and not well described by a semiclassical space-time) - just enough to let the single qubit Q pass through. Therefore, we propose that the gravitational dual description of quantum teleportation understood as a dynamical process is that the qubit passes through the ER=EPR wormhole of the entangled pair, A and B , which has been rendered traversable by the required interaction.

Another possible interpretation of our result is to relate it to the recovery of information described in [24]. Assuming that black hole evaporation is unitary, it is in principle possible to eventually recover a qubit which falls into a black hole, from a quantum computation

acting on the Hawking radiation. Assuming that you have access to an auxiliary system maximally entangled with the black hole, and that the black hole is an efficient scrambler of information, it turns out that you only need a small (order unity) additional quantity of Hawking radiation to reconstruct the qubit. In our system, the qubit may be identified with the system that falls into the black hole from the left and gets scrambled, the auxiliary entangled system is the CFT on the right, and the boundary interaction somehow triggers the appropriate quantum computation to make the qubit reappear again, after a time of order the scrambling time $R \ln(R/L_{\text{planck}})$.⁹

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A $\int dUT_{UU}$

Using (3.8), the integrated null energy is

$$\int_{U_0}^{\infty} dUT_{UU} = -4h\Delta C_0 \int_{U_0}^{\infty} dU \lim_{U' \rightarrow U} \partial_U G(U, U'; U_0), \quad (\text{A.1})$$

where

$$G(U, U'; U_0) \equiv \int_{U_0}^U dU_1 \int_1^{\frac{U}{U_1}} \frac{dy}{\sqrt{y^2 - 1}} \frac{U_1^\Delta}{(U - U_1 y)^\Delta (U' U_1 + y)^{\Delta+1}}. \quad (\text{A.2})$$

Note that

$$\lim_{U' \rightarrow U} \partial_U G(U, U'; U_0) = \partial_U G(U, U; U_0) - \partial_U^{(2)} G(U, U; U_0), \quad (\text{A.3})$$

where $\partial_U^{(2)}$ indicates a derivative with respect to the second variable. (A.1) becomes

$$\int_{U_0}^{\infty} dUT_{UU} = -4h\Delta C_0 \left(G(\infty, \infty; U_0) - G(U_0, U_0; U_0) - \int_{U_0}^{\infty} dU \partial_U^{(2)} G(U, U; U_0) \right) \quad (\text{A.4})$$

By (3.9), and changing to the z variable,

$$\begin{aligned} G(U, U; U_0) &\propto \int_{U_0}^U dU_1 \frac{U_1^{\Delta-1/2}}{(U - U_1)^{\Delta-1/2} (1 + UU_1)^{\Delta+1}} F_1 \left(\frac{1}{2}; \frac{1}{2}, \Delta + 1; \frac{3}{2} - \Delta; \frac{U_1 - U}{2U_1}, \frac{U_1 - U}{U_1(1 + UU_1)} \right) \\ &= \int_0^1 dz \frac{((U - U_0)z + U_0)^{\Delta-1/2} F_1 \left(\frac{1}{2}; \frac{1}{2}, \Delta + 1; \frac{3}{2} - \Delta; -\frac{(U - U_0)(1-z)}{2((U - U_0)z + U_0)}, -\frac{(U - U_0)(1-z)}{((U - U_0)z + U_0)(1 + U((U - U_0)z + U_0))} \right)}{(U - U_0)^{\Delta-3/2} (1 - z)^{\Delta-1/2} (1 + U((U - U_0)z + U_0))^{\Delta+1}} \end{aligned} \quad (\text{A.5})$$

⁹We thank Juan Maldacena for suggesting this interpretation.

which immediately implies $G(U_0, U_0; U_0) = 0$ given that $\Delta < 3/2$. For the large U limit, $G(\infty, \infty; U_0)$, the prefactor of F_1 above decays at least as fast as $U^{-\Delta}$ and the F_1 part becomes

$$\begin{aligned} F_1\left(\frac{1}{2}; \frac{1}{2}; \Delta + 1; \frac{3}{2} - \Delta; -\frac{1-z}{2z}, 0\right) &= {}_2F_1\left(\frac{1}{2}, \frac{1}{2}; \frac{3}{2} - \Delta; \frac{z-1}{2z}\right) \\ &= \left(\frac{2z}{z+1}\right)^{1/2} {}_2F_1\left(\frac{1}{2}, \frac{1}{2}; \frac{3}{2} - \Delta; \frac{1-z}{1+z}\right) \end{aligned} \quad (\text{A.6})$$

which leads to

$$G(\infty, \infty; U_0) \sim U^{-\Delta} \int_0^1 dz \left(\frac{2z}{z+1}\right)^{1/2} (1-z)^{-\Delta+1/2} {}_2F_1\left(\frac{1}{2}, \frac{1}{2}; \frac{3}{2} - \Delta; \frac{1-z}{1+z}\right) \rightarrow 0 \quad (\text{A.7})$$

where in the last step we used the fact that the z integral is finite due to the property of hypergeometric function:

$${}_2F_1\left(\frac{1}{2}, \frac{1}{2}; \frac{3}{2} - \Delta; 0\right) = 1, \quad \lim_{z \rightarrow 0} z^{1/2} {}_2F_1\left(\frac{1}{2}, \frac{1}{2}; \frac{3}{2} - \Delta; \frac{1-z}{1+z}\right) \sim z^{1/2} \log \frac{2z}{z+1} \rightarrow 0 \quad (\text{A.8})$$

The integral of T_{UU} is simplified to be

$$\begin{aligned} & -\frac{1}{4h\Delta C_0} \int_{U_0}^{\infty} dU T_{UU} \\ &= -\int_{U_0}^{\infty} dU \int_{U_0}^U dU_1 \int_1^{\frac{U}{U_1}} \lim_{U' \rightarrow U} \partial_{U'} \frac{dy}{\sqrt{y^2-1}} \frac{U_1^\Delta}{(U-U_1y)^\Delta (U'U_1+y)^{\Delta+1}} \\ &= \int_{U_0}^{\infty} dU_1 \int_{U_1}^{\infty} dU \int_1^{\frac{U}{U_1}} \frac{dy}{\sqrt{y^2-1}} \frac{(\Delta+1)U_1^{\Delta+1}}{(U-U_1y)^\Delta (UU_1+y)^{\Delta+2}} \\ &= \int_{U_0}^{\infty} dU_1 \int_{U_1}^{\infty} dU \frac{(\Delta+1)\Gamma(\frac{1}{2})\Gamma(1-\Delta)U_1^{2\Delta+3}(U+U_1)^{-1/2}}{\Gamma(\frac{3}{2}-\Delta)(U-U_1)^{\Delta-1/2}U^{\Delta+2}(1+U_1^2)^{\Delta+2}} \\ & \quad \times F_1\left(1-\Delta; \frac{1}{2}; \Delta+2; \frac{3}{2}-\Delta; \frac{U-U_1}{U+U_1}, \frac{U-U_1}{U(1+U_1^2)}\right) \end{aligned} \quad (\text{A.9})$$

For further simplification, we use (3.17) and define $w = \frac{U-U_1}{U+U_1}$ to rewrite (A.9) as

$$\begin{aligned} & -\frac{1}{4h\Delta C_0} \int_{U_0}^{\infty} dU T_{UU} \\ &= \sum_m \frac{(\Delta+1)\Gamma(\frac{1}{2})\Gamma(1-\Delta)(1-\Delta)_m(\Delta+2)_m 2^{m+1-\Delta}}{m!\Gamma(\frac{3}{2}-\Delta)(\frac{3}{2}-\Delta)_m} \int_{U_0}^{\infty} dU_1 \frac{U_1^2}{(1+U_1^2)^{\Delta+m+2}} \\ & \quad \times \int_0^1 dw \frac{w^{m-\Delta+1/2}(1-w)^{2\Delta}}{(1+w)^{\Delta+m+2}} {}_2F_1\left(1-\Delta+m; \frac{1}{2}; \frac{3}{2}-\Delta+m; w\right) \\ &= \sum_m \frac{(\Delta+1)\Gamma(\frac{1}{2})\Gamma(1-\Delta)(1-\Delta)_m(\Delta+2)_m 2^{m+1-\Delta}}{m!\Gamma(\frac{3}{2}-\Delta)(\frac{3}{2}-\Delta)_m} \int_{U_0}^{\infty} dU_1 \frac{U_1^2}{(1+U_1^2)^{\Delta+m+2}} \\ & \quad \times \frac{\Gamma(\frac{3}{2}-\Delta+m)\Gamma(2\Delta+1)^2}{2^{\Delta+m+2}\Gamma(\frac{3}{2}+2\Delta)\Gamma(2+\Delta+m)} {}_2F_1\left(2\Delta+1, 2\Delta+1; \frac{3}{2}+2\Delta; \frac{1}{2}\right) \end{aligned}$$

$$\begin{aligned}
&= \sum_m \frac{\Gamma(\frac{1}{2})\Gamma(1-\Delta)\Gamma(2\Delta+1)^2(1-\Delta)_m}{m!\Gamma(\frac{3}{2}+2\Delta)\Gamma(\Delta+1)2^{2\Delta+1}} {}_2F_1\left(2\Delta+1, 2\Delta+1; \frac{3}{2}+2\Delta; \frac{1}{2}\right) \\
&\quad \times \frac{(U_0^{-2})^{\Delta+m+1/2}}{2(\Delta+m+1/2)} {}_2F_1\left(\frac{1}{2}+\Delta+m, 2+\Delta+m; \frac{3}{2}+\Delta+m; -U_0^{-2}\right) \\
&= \frac{\Gamma(\frac{1}{2})\Gamma(1-\Delta)\Gamma(2\Delta+1)^2}{2^{2\Delta+2}(\Delta+\frac{1}{2})\Gamma(\frac{3}{2}+2\Delta)\Gamma(\Delta+1)} \frac{1}{(1+U_0^2)^{\Delta+1/2}} {}_2F_1\left(2\Delta+1, 2\Delta+1; \frac{3}{2}+2\Delta; \frac{1}{2}\right) \\
&\quad \times \sum_m \frac{(1-\Delta)_m(\frac{1}{2}+\Delta)_m}{m!(\frac{3}{2}+\Delta)_m} \left(\frac{1}{1+U_0^2}\right)^m {}_2F_1\left(\frac{1}{2}+\Delta+m, -\frac{1}{2}; \frac{3}{2}+\Delta+m; \frac{1}{1+U_0^2}\right) \\
&= \frac{\pi\Gamma(1-\Delta)\Gamma(2\Delta+1)^2}{2^{2\Delta+2}(\Delta+\frac{1}{2})\Gamma(\Delta+1)^3} \frac{{}_2F_1\left(\frac{1}{2}+\Delta, \frac{1}{2}-\Delta; \frac{3}{2}+\Delta; \frac{1}{1+U_0^2}\right)}{(1+U_0^2)^{\Delta+1/2}} \tag{A.10}
\end{aligned}$$

where in fifth line we used [43]

$$\begin{aligned}
&\int_0^y \frac{x^{c-1}(y-x)^{\beta-1}}{(1-zx)^\rho} {}_2F_1(a, b; c; \frac{x}{y}) dx \\
&= \frac{y^{c+\beta-1}}{(1-yz)^\rho} \frac{\Gamma(c)\Gamma(\beta)\Gamma(c-a-b+\beta)}{\Gamma(c-a+\beta)\Gamma(c-b+\beta)} {}_3F_2\left(\beta, \rho, c-a-b+\beta; c-a+\beta, c-b+\beta; \frac{yz}{yz-1}\right) \\
&\quad [y, \Re c, \Re \beta, \Re(c-a-b+\beta) > 0; |\arg(1-yz)| < \pi] \tag{A.11}
\end{aligned}$$

in sixth line we used

$$\int_b^\infty dx \frac{x^2}{(1+x^2)^a} = \frac{b^{-2a+3}}{2a-3} {}_2F_1\left(a-\frac{3}{2}, a; a-\frac{1}{2}; -b^{-2}\right) \tag{A.12}$$

and in last step we used [43]

$$\sum_{k=0}^\infty \frac{(a)_k(b')_k}{k!(c)_k} x^k {}_2F_1(a+k, b; c+k; x) = {}_2F_1(a, b+b'; c; x) \tag{A.13}$$

$${}_2F_1\left(2\Delta+1, 2\Delta+1; \frac{3}{2}+2\Delta; \frac{1}{2}\right) = \frac{\pi^{1/2}\Gamma(\frac{3}{2}+2\Delta)}{\Gamma(1+\Delta)^2} \tag{A.14}$$

In the end, restoring ℓ , we find the following relatively simple result

$$\int_{U_0}^\infty dUT_{UV} = -\frac{h\Gamma(2\Delta+1)^2}{2^{4\Delta}(2\Delta+1)\Gamma(\Delta)^2\Gamma(\Delta+1)^2\ell} \frac{{}_2F_1\left(\frac{1}{2}+\Delta, \frac{1}{2}-\Delta; \frac{3}{2}+\Delta; \frac{1}{1+U_0^2}\right)}{(1+U_0^2)^{\Delta+1/2}} \tag{A.15}$$

If we turn off the interaction at U_f , the integral is just the difference between $\int_{U_0}^\infty dUT_{UV}$ and $\int_{U_f}^\infty dUT_{UV}$.

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